TIDAL DISSIPATION IN RUBBLE-PILE BINARY ASTEROIDS. F. Nimmo[1], I. Matsuyama[2] [1]Dept. Earth and Planetary Sciences, University of California Santa Cruz, Santa Cruz, CA 95064, USA [2] Lunar and Planetary Laboratory, University of Arizona, Tucson, AZ 85721, USA

Summary: We derive a simple theory describing tidal dissipation in the regolith layer of rubble-pile binary asteroids [1]. The theory agrees with inferred tidal dissipation rates if the regolith thickness is independent of body size. Applications to wobble, librations and eccentricity/semi-major axis excitation are discussed.

Introduction: The tidal response of a body provides information on its internal structure and mechanical properties [2]. The rate of tidal dissipation depends on k_2/Q or alternatively $1/\mu Q$, where k_2 is the tidal Love number (giving the response amplitude), Q is the dissipation factor (giving the response phase) and μ is the effective rigidity [3,4].

Observations: Jacobson and Scheeres [5] used binary asteroids to determine k_2/Q, by assuming that the observed semi-major axis was a result of equilibrium between tidal dissipation and the binary YORP effect. An alternative [6] is to neglect the second effect and instead assume a system age; this yields a higher (less dissipative) k_2/Q value. A recent astrometric study of 1996FG$_3$ [7] showed no semi-major axis evolution and thus supports the equilibrium assumption.

Theory: An important theoretical treatment by [8] showed that yielding in rubble-pile asteroids results in an effective rigidity much less than that of a monolithic body, and predicts that k_2 scales with primary radius R.

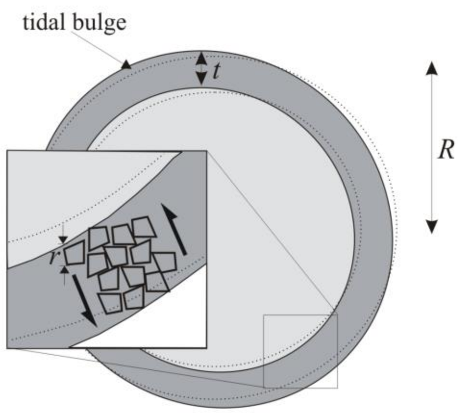

Figure 1: Sketch of the geometry of the problem. Tidal deformation is represented by the departure of the solid surface from the mean shape (dashed lines) and results in shear strains (indicated by half-arrows).

We focus on tidal dissipation in a regolith layer of thickness t (Figure 1). The motion of the tidal bulge relative to the surface results in shearing motion of individual regolith blocks. On a single block face the frictional dissipation rate depends on the shear stress, the surface area and the sliding velocity, and is given by

$$\dot{E}_f \sim (\rho g t) f r^2 u \Omega_p \qquad (1)$$

where ρ is the density, g the surface acceleration, r the block size, u the relative displacement between neighbouring faces, Ω_p the spin rate and f the friction coefficient. The displacement depends on the tidal strain ε and the block dimensions and may be written

$$u \sim \epsilon r \sim r\left[\frac{H}{R}\right] \sim r h_2 \left[\frac{q}{1+q}\frac{n^2}{G\rho}\right] \qquad (2)$$

where H is the amplitude of the tidal bulge, h_2 is the tidal displacement Love number, q is the mass ratio between secondary and primary, n is the mean motion of the secondary and G is the gravitational constant.

Combining equations (1) and (2) we derive the total dissipation rate in the regolith layer of

$$\dot{E}_f \sim N f h_2 M q n^2 \Omega_p t^2 \qquad (3)$$

Here N (~3 for a roughly cubic element) is the number of faces per element, m is the mass of the secondary and we have dropped the (1+q) term which is generally close to unity. The t^2 term arises because increasing t increases the overburden pressure and the total dissipative volume. As expected, dissipation increases with friction coefficient, forcing frequency and displacement (h_2); it is also independent of the element size r, as long as $r<<t$.

Equation (3) may be compared with the conventional expression for tidal dissipation in a non-synchronous body [9] to derive an effective Q, given by

$$Q_{eff} \sim \left[\frac{qn^2}{G\rho}\right]\frac{1}{Nf}\left(\frac{R}{t}\right)^2 \sim 3\times10^{-3}\left(\frac{qn^2}{3\times10^{-10}\ s^{-2}}\right)\left(\frac{2\ g/cc}{\rho}\right)\left(\frac{R}{t}\right)^2$$

where here we have assumed that $h_2 \approx k_2$.

This result may then be combined with the prediction of [8] to derive equation (4):

$$\frac{Q}{k_2} \sim 3\times10^5\left(\frac{R}{1\ km}\right)\left(\frac{qn^2}{3\times10^{-10}\ s^{-2}}\right)\left(\frac{2\ g/cc}{\rho}\right)\left(\frac{30\ m}{t}\right)^2$$

The same result can also be used to predict the quantity μQ, here given in SI units:

$$\mu Q \sim 10^8\left(\frac{R}{1\ km}\right)^3\left(\frac{qn^2}{3\times10^{-10}}\right)\left(\frac{30\ m}{t}\right)^2 \qquad (5)$$

Comparison with observations: We use the approach of [5] but with an expanded catalogue of asteroid binaries, from [10] and using a BYORP parameter

B=10^{-2} [7]. The inferred Q/k_2 as function of body radius is shown in Fig 2. As noted by [5], the inferred Q/k_2 scales roughly with R (or $R^{1.5}$), which is opposite to the prediction of [8] if Q is constant. In contrast, the observations are consistent with equation (4) if t is constant, or decreases slightly with radius. Furthermore, the dependence on qn^2, shown by colours in Figure 2, is also approximately consistent with equation (4)

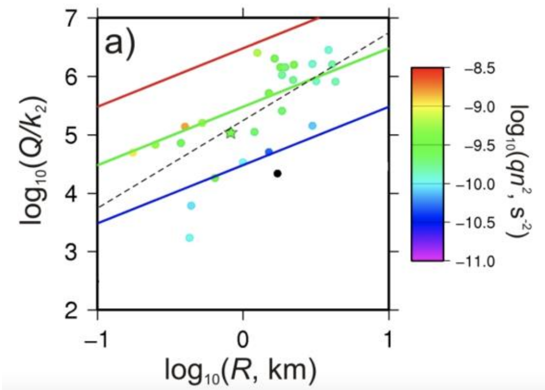

Figure 2: Dots are data plotted from [1] taking $B = 10^{-2}$ (see text); colour indicates the quantity qn^2. Star is (175706) 1996 FG$_3$ [7]. Dashed line shows least-squares fit to the data, with a gradient of 1.51. Coloured lines use equation 4 with three different values of qn^2 ($10^{-8.5}$, $10^{-9.5}$, $10^{-10.5}$ s^{-2}) and t=30 m.

A further observation of relevance is a study of tumbling asteroids by [11]. These authors argue that the damping timescale for such tumbling is approximately independent of radius. For this to be the case, μQ would need to scale as R^2. Our simple analysis (equation 5) predicts an R^3 dependence; thus, there is qualitative but not quantitative agreement.

Regolith Thickness: Figure 2 suggests that the regolith thickness t~30m, independent of body radius. Only rather scanty estimates of regolith thickness are available: a few metres or more on Itokawa (R=0.17 km) [12]; 30-200 m on Gaspra (R=6.1 km) [13]; up to a few tens of metres on Eros (R=8.4 km) [14]; 100-200 m on Phobos (R=11.3 km) [15]; ~50m on Ida (R=15.7 km) [16]. Based on these results it certainly seems as if regolith thickness varies only rather weakly with radius, and a ~30m thickness would be hard to rule out. A theoretical study by [17] argued that the expected regolith thickness is tens of metres, and should decrease slightly with R. This prediction is in good agreement with our results.

Application to wobble: For an isolated rubble-pile asteroid undergoing a wobble of amplitude α the dissipation rate within a regolith layer is

$$\dot{E_f} \sim N f h_2 M \alpha \Omega_w {\Omega_p}^2 t^2 \qquad (6)$$

which is analogous to equation (3) with the strain rate determined by the rotational bulge, the wobble angular frequency Ω_w and amplitude α. The quantity Ω_w is smaller than Ω_p by a factor of $h_2\Omega_p^2/G\rho$ [18]. Using the total wobble energy from [18], the damping timescale may be written

$$\tau_w \sim \left(\frac{R}{t}\right)^2 \frac{\alpha}{N f h_2}\frac{1}{\Omega_p} \qquad (7)$$

The damping timescale decreases with increasing friction or bulge amplitude (h_2) as expected. For a Bennu-size asteroid, R~300m implies h_2~3x10^{-6} [8]. For a regolith thickness of 30m the damping timescale is then ~10^6 rotation periods, or ~10^3 years. Spacecraft missions should be capable of measuring such damping. The damping timescale is shorter than conventional estimates [18], primarily because h_2 is larger than the equivalent calculation for a monolithic body.

Eccentricity & semi-major axis excitation: In BYORP equilibrium the eccentricity of the secondary is expected to be constant [5]. However, for binaries not in equilibrium, dissipation in the primary is expected to dominate [19], which will excite the secondary's eccentricity and increase its semi-major axis. The characteristic timescale for both processes is

$$\tau_{a,e} \sim 4\ kyr \left(\frac{a}{R}\right)^{7/2} \left(\frac{R}{1\ km}\right) \left(\frac{30m}{t}\right)^2 \qquad (8)$$

where a is the semi-major axis. For a=3R this timescale is ~0.3 Myr, short compared to solar system timescales and potentially measurable via astrometry.

Librations: Dissipation will change the amplitude and phase of the forced librations of body, such as Phobos. However, the phase lag depends on k_2/Q [20,21] and will thus be too small to measure directly. In contrast, damping of excited free librations (e.g. the Janus-Epimetheus pair [22]) for rubble-pile bodies is expected to be rapid and potentially measurable.

References: [1] Nimmo & Matsuyama, *Icarus* 2019 [2] Moore & Schubert, *Icarus* 2000 [3] Munk & MacDonald, *The rotation of the Earth*, 1960 [4] Efroimsky & Makarov, *Ap.J.* 2013 [5] Jacobson & Scheeres, *Ap. J. L.* 2011 [6] Taylor & Margot, *Icarus* 2011 [7] Scheirich et al., *Icarus* 2015 [8] Goldreich & Sari, *Ap. J.* 2009 [9] Murray & Dermott, *Solar System Dynamics*, 1999 [10] Pravec et al., *Icarus* 2016 [11] Pravec et al. , *Icarus* 2014 [12] Barnouin-Jha et al., *Icarus* 2008 [13] Veverka et al., *Icarus* 1994 [14] Prockter et al., *Icarus* 2002 [15] Thomas et al., *JGR* 1979 [16] Sullivan et al., *Icarus* 1996 [17] Langevin & Maurette, *LPSC* 1980 [18] Burns & Safronov, *MNRAS*, 1973 [19] Goldreich, *MNRAS*, 1963. [20] Rambaux et al., *GRL* 2010. [21] Caudal, *Icarus* 2017. [22] Tiscareno et al., *Icarus* 2009.